BLACK HOLE PROBES IN STABILIZED VACUA

 $\phi(r)$

 r_H

ifl

Ø

()

Matilda Delgado

Based on: [24xx.xxxx] M. Delgado, D. Mayerson, S. Reymond, T. Van Riet

BIG PICTURE

Running solutions are solutions to the EOMs where scalar fields depend on spacetime coordinates



BIG PICTURE

Running solutions are solutions to the EOMs where scalar fields depend on spacetime coordinates



For example:

-Any scalar field with a **run-away potential** -the backreaction of **branes** that sources the scalar field

[...]



They are ubiquitous in string theory (especially once you break SUSY)

 $\phi(r$

 r_H

BIG PICTURE

Running solutions are solutions to the EOMs where scalar fields depend on spacetime coordinates



For example:

-Any scalar field with a **run-away potential** -the backreaction of **branes** that sources the scalar field

[...]



They provide a map between spacetime and moduli space:

→ they can be used as 'experimental' probes of different corners of moduli space
 → in particular, the boundaries of moduli space where UV physics manifest

Specific running solution: charged Black Holes in stringy EFTs

Specific running solution: charged Black Holes in stringy EFTs

Since the U(1) gauge couplings are functions of scalars, charged BHs create a **potential** for these fields

There is an **attractor mechanism** where the value of the scalar fields at the horizon is entirely determined by the black hole's charges



Specific running solution: charged Black Holes in stringy EFTs

Since the U(1) gauge couplings are functions of scalars, charged BHs create a **potential** for these fields

There is an **attractor mechanism** where the value of the scalar fields at the horizon is entirely determined by the black hole's charges

Celebrated example: the 4d N=2 attractor mechanism [S. Ferrara, R. Kallosh, A. Strominger '95]



Specific running solution: charged Black Holes in stringy EFTs

Since the U(1) gauge couplings are functions of scalars, charged BHs create a **potential** for these fields

There is an **attractor mechanism** where the value of the scalar fields at the horizon is entirely determined by the black hole's charges

These large (smooth) BHs were used to probe infinite distances limits in moduli space For example:

Specific running solution: charged Black Holes in stringy EFTs

Since the U(1) gauge couplings are functions of scalars, charged BHs create a potential for these fields

There is an **attractor mechanism** where the value of the scalar fields at the horizon is entirely determined by the black hole's charges

These large (smooth) BHs were used to probe infinite distances limits in moduli space

For example:

-> The BH entropy distance conjecture

[Q. Bonnefoy, L. Ciambelli, D. Lüst, S. Lüst '19]

Specific running solution: charged Black Holes in stringy EFTs

Since the U(1) gauge couplings are functions of scalars, charged BHs create a potential for these fields

There is an **attractor mechanism** where the value of the scalar fields at the horizon is entirely determined by the black hole's charges

These large (smooth) BHs were used to probe infinite distances limits in moduli space

For example:

-> The BH entropy distance conjecture

-> Obtain topological data about underlying compact space from BH thermodynamics

[Q. Bonnefoy, L. Ciambelli, D. Lüst, S. Lüst '19]

[Delgado, Montero, Vafa '22]

Specific running solution: charged Black Holes in stringy EFTs

Since the U(1) gauge couplings are functions of scalars, charged BHs create a **potential** for these fields

There is an **attractor mechanism** where the value of the scalar fields at the horizon is entirely determined by the black hole's charges

These large (smooth) BHs were used to probe infinite distances limits in moduli space

BUT ALL OF THIS WAS DONE IN TRUE MODULI SPACES

Specific running solution: charged Black Holes in stringy EFTs

Since the U(1) gauge couplings are functions of scalars, charged BHs create a **potential** for these fields

There is an **attractor mechanism** where the value of the scalar fields at the horizon is entirely determined by the black hole's charges

These large (smooth) BHs were used to probe infinite distances limits in moduli space

BUT ALL OF THIS WAS DONE IN TRUE MODULI SPACES

Today: how does being in a scale-separated stabilized vacuum change the story?

Today: how does being in a scale-separated stabilized vacuum change the story?

Why do we expect anything to change?

Today: how does being in a scale-separated stabilized vacuum change the story?

Why do we expect anything to change?

Simply put, there will be a competition between the stabilizing potential and the black hole potential.

The BH can break the vacuum OR the stabilizing potential can change/break the attractor mechanism [D. Green, E. Silverstein, D. Starr '06]

Similar phenomena were also observed in [R Anglus, J. Huertas, A. Uranga, '23] for the case of run-away potentials.

Today: how does being in a scale-separated stabilized vacuum change the story?

Why do we expect anything to change?

Simply put, there will be a **competition** between the **stabilizing potential** and the **black hole potential**.

The BH can break the vacuum OR the stabilizing potential can change/break the attractor mechanism [D. Green, E. Silverstein, D. Starr '06]

Similar phenomena were also observed in [R Angius, J. Huertas, A. Uranga, '23] for the case of run-away potentials.

Understanding this can teach us:

Today: how does being in a scale-separated stabilized vacuum change the story?

Why do we expect anything to change?

Simply put, there will be a **competition** between the **stabilizing potential** and the **black hole potential**.

The BH can break the vacuum OR the stabilizing potential can change/break the attractor mechanism [D. Green, E. Silverstein, D. Starr '06]

Similar phenomena were also observed in [R Anglus, J. Huertas, A. Uranga, '23] for the case of run-away potentials.

Understanding this can teach us:

↔ how to use these probe BHs in more realistic phenomenological (?) scenarios

Today: how does being in a scale-separated stabilized vacuum change the story?

Why do we expect anything to change?

Simply put, there will be a **competition** between the **stabilizing potential** and the **black hole potential**.

The BH can break the vacuum OR the stabilizing potential can change/break the attractor mechanism [D. Green, E. Silverstein, D. Starr '06]

Similar phenomena were also observed in [R Anglus, J. Huertas, A. Uranga, '23] for the case of run-away potentials.

Understanding this can teach us:

→ how to use these **probe BHs** in more realistic phenomenological (?) scenarios

⇒ about scale-separation in QG in general

Today: how does being in a scale-separated stabilized vacuum change the story?

Why do we expect anything to change?

Simply put, there will be a **competition** between the **stabilizing potential** and the **black hole potential**.

The BH can break the vacuum OR the stabilizing potential can change/break the attractor mechanism [D. Green, E. Silverstein, D. Starr '06]

Similar phenomena were also observed in [R Anglus, J. Huertas, A. Uranga, '23] for the case of run-away potentials.

Understanding this can teach us:

↔ how to use these probe BHs in more realistic phenomenological (?) scenarios

⇒ about scale-separation in QG in general

.... Let's start with a toy model

Dilatonic Charged Black Hole in the presence of a dilaton mass:

$$S = \int dx^4 \sqrt{-g} \left\{ R - \frac{1}{2} (\partial \phi)^2 - \frac{1}{4} e^{a\phi} F^2 - \frac{1}{2} m^2 \phi^2 \right\}$$

Dilatonic Charged Black Hole in the presence of a dilaton mass:

$$S = \int dx^4 \sqrt{-g} \left\{ R - \frac{1}{2} (\partial \phi)^2 - \frac{1}{4} e^{a\phi} F^2 - \frac{1}{2} m^2 \phi^2 \right\}$$

1) Away from the BH, only the mass contributes and the dilaton is stabilized to zero Typical scale: m

φ

 $V(\phi)$

Dilatonic Charged Black Hole in the presence of a dilaton mass:

$$S = \int dx^4 \sqrt{-g} \left\{ R - \frac{1}{2} (\partial \phi)^2 - \frac{1}{4} e^{a\phi} F^2 - \frac{1}{2} m^2 \phi^2 \right\}$$

1) Away from the BH, only the mass contributes and the dilaton is stabilized to zero Typical scale: m

Φ

 $V(\phi)$

2) Close to the BH, the BH also contributes and tries to attract the dilaton to a new value at the horizon Typical scale: R(P,Q)



To figure out what happens near the BH, consider the set of equations:



[M. Montero, T. Van Riet, and G. Venken, '20]

To figure out what happens near the BH, consider the set of equations:

lf mR << 1 :

The **mass of the dilaton is negligible**, and the dilaton is effectively a proper modulus

The black hole reduces to the dilatonic BH solution.

(This is the usual set-up in the literature)



[M. Montero, T. Van Riet, and G. Venken, '20]

24

A TOY MODEL

To figure out what happens near the BH, consider the set of equations:

<u>|fmR<<1:</u>

The **mass of the dilaton is negligible**, and the dilaton is effectively a proper modulus

The black hole reduces to the dilatonic BH solution.

(This is the usual set-up in the literature)



lf mR >> 1:

25/06/24

The **stabilizing potential dominates** and the dilaton is forced to stay at its VEV

The black hole becomes a simple Reissner-Nordström charged BH.

M. Montero, T. Van Riet, and G. Venken, '20]

25/06/24

A TOY MODEL

To figure out what happens near the BH, consider the set of equations:

<u>|fmR<<1:</u>

The **mass of the dilaton is negligible**, and the dilaton is effectively a proper modulus

The black hole reduces to the dilatonic BH solution.

(This is the usual set-up in the literature)



<u>lf mR >> 1 :</u>

The **stabilizing potential dominates** and the dilaton is forced to stay at its VEV

The black hole becomes a simple Reissner-Nordström charged BH.

[M. Montero, T. Van Riet, and G. Venken, '20]

25/06/24

26

A TOY MODEL

To figure out what happens near the BH, consider the set of equations:

<u>|fmR<<1:</u>

The **mass of the dilaton is negligible**, and the dilaton is effectively a proper modulus

The black hole reduces to the dilatonic BH solution.

(This is the usual set-up in the literature)



[M. Montero, T. Van Riet, and G. Venken, '20]

<u>lf mR >> 1 :</u>

The **stabilizing potential dominates** and the dilaton is forced to stay at its VEV

The black hole becomes a simple Reissner-Nordström charged BH.

Let's see how the RN BH is modified is we make the dilaton a little lighter

Consider the limit where the dilaton is extremely massive and do perturbation theory in:

$$= (m^2 R^2)^{-1} << 1$$

Consider the limit where the dilaton is extremely massive and do perturbation theory in:

$$x = (m^2 R^2)^{-1} <<$$

At 0-th order, we have the following at the horizon:

$$egin{cases} \phi_0=0 & ext{Massive dilaton stabilized at 0} \ R_0^2=P^2+Q^2 & ext{RN Black Hole} \end{cases}$$

Consider the limit where the dilaton is extremely massive and do perturbation theory in:

$$\epsilon = (m^2 R^2)^{-1} <<$$

At 0-th order, we have the following at the horizon:

Including the first correction, we have:

$$\left\{ egin{array}{ll} \phi_0 = 0 & {
m Massive} \ {
m dilaton} \ {
m stabilized} \ {
m at} \ 0 & {
m R}_0^2 = P^2 + Q^2 & {
m RN} \ {
m Black} \ {
m Hole} \end{array}
ight.$$

$$\begin{cases} \phi = 0 - 2a \frac{R_0^4}{r^4} \frac{P^2 - Q^2}{P^2 + Q^2} \epsilon + \mathcal{O}(\epsilon^2) \\ R^2 = R_0^2 \left(1 - a^2 \left(\frac{P^2 - Q^2}{P^2 + Q^2} \right)^2 \epsilon + \mathcal{O}(\epsilon^2) \right) \end{cases}$$

Consider the limit where the dilaton is extremely massive and do perturbation theory in:

$$\epsilon = (m^2 R^2)^{-1} <<$$

At 0-th order, we have the following at the horizon:

Including the first correction, we have:

$$\left\{ egin{array}{ll} \phi_0 = 0 & {
m Massive dilaton stabilized at 0} \ R_0^2 = P^2 + Q^2 & {
m RN Black Hole} \end{array}
ight.$$

$$\begin{cases} \phi = 0 - 2a \frac{R_0^4}{r^4} \frac{P^2 - Q^2}{P^2 + Q^2} \epsilon + \mathcal{O}(\epsilon^2) \\ R^2 = R_0^2 \left(1 - a^2 \left(\frac{P^2 - Q^2}{P^2 + Q^2} \right)^2 \epsilon + \mathcal{O}(\epsilon^2) \right) \end{cases}$$

➡ the dilaton can never blow up no matter what charges you pick
Exploring infinite distances in moduli space seems impossible

Consider the limit where the dilaton is extremely massive and do perturbation theory in:

$$\epsilon = (m^2 R^2)^{-1} <<$$

At 0-th order, we have the following at the horizon:

Including the first correction, we have:

$$egin{pmatrix} \phi_0 = 0 & ext{Massive dilaton stabilized at 0} \ R_0^2 = P^2 + Q^2 & ext{RN Black Hole} \ \end{pmatrix}$$

$$\begin{cases} \phi = 0 - 2a \frac{R_0^4}{r^4} \frac{P^2 - Q^2}{P^2 + Q^2} \epsilon + \mathcal{O}(\epsilon^2) \\ R^2 = R_0^2 \left(1 - a^2 \left(\frac{P^2 - Q^2}{P^2 + Q^2} \right)^2 \epsilon + \mathcal{O}(\epsilon^2) \right) \end{cases}$$

➡ the dilaton can never blow up no matter what charges you pick
Exploring infinite distances in moduli space seems impossible

➡ Probing farther points in moduli space corresponds to taking a smaller (more singular) black hole

I KNOW WHAT YOU'RE THINKING...





The scalar potential we have chosen,

$$V(\phi) = \frac{1}{2}m^2\phi^2$$

is probably in the Swampland.

I KNOW WHAT YOU'RE THINKING..



The scalar potential we have chosen,

$$V(\phi) = \frac{1}{2}m^2\phi^2$$

is probably in the Swampland.

How can we know that our toy model would actually describe what happens in string theory?

The next step has to be:

35

I KNOW WHAT YOU'RE THINKING..



The scalar potential we have chosen,

$$V(\phi) = \frac{1}{2}m^2\phi^2$$

is probably in the Swampland.

How can we know that our toy model would actually describe what happens in string theory?

The next step has to be:

a top-down example!

We need to build a black hole in a stabilized, scale separated vacuum:

 $L_{KK} << L_{c.c.}$

This is easier said than done,

We need to build a black hole in a stabilized, scale separated vacuum:

$$L_{KK} << L_{c.c.}$$

This is easier said than done,

In DGKT [O. DeWolfe, A. Giryavets, S. Kachru, W. Taylor '05], we can get a scale-separated AdS vacuum

Explicit scenario: A toroidal Orbifold Model $\, \mathbb{T}^{6}/\mathbb{Z}_{4} \,$

[M. Ihl, (D. Robbins), T. Wrase, "06, '07]

Simple because it has 1 U(1) whose gauge kinetic function depends on 1 single modulus

We need to build a black hole in a stabilized, scale separated vacuum:

$$L_{KK} << L_{c.c.}$$

This is easier said than done,

In DGKT [O. DeWolfe, A. Giryavets, S. Kachru, W. Taylor '05], we can get a scale-separated AdS vacuum

Explicit scenario: A toroidal Orbifold Model $\, \mathbb{T}^{6}/\mathbb{Z}_{4} \,$

[M. Ihl, (D. Robbins), T. Wrase, "06, '07]

Simple because it has 1 U(1) whose gauge kinetic function depends on 1 single modulus

Impossible to do perturbation theory (like in the toy model) since moduli have masses at the AdS scale

C.f. AdS moduli conjecture [F.F. Gautason, V. Van Hemelryck, T. Van Riet '18]

We need to build a black hole in a stabilized, scale separated vacuum:

$$L_{KK} << L_{c.c.}$$

This is easier said than done,

In DGKT [O. DeWolfe, A. Giryavets, S. Kachru, W. Taylor '05], we can get a scale-separated AdS vacuum

Explicit scenario: A toroidal Orbifold Model $\, \mathbb{T}^{6}/\mathbb{Z}_{4} \,$

[M. Ihl, (D. Robbins), T. Wrase, "06, '07]

Simple because it has 1 U(1) whose gauge kinetic function depends on 1 single modulus

Impossible to do perturbation theory (like in the toy model) since moduli have masses at the AdS scale C.f. AdS moduli conjecture [F.F. Gautason, V. Van Hemelryck, T. Van Riet '18]

Numerical approach?



We need to build a black hole in a stabilized, scale separated vacuum:

In GKP [S. Giddings, S. Kachru, J. Polchinski '01], we can get a Minkowski vacuum

but there are U(1)'s that only couple to the complex structure (stabilized) moduli,

We need to build a black hole in a stabilized, scale separated vacuum:

In GKP [S. Giddings, S. Kachru, J. Polchinski '01], we can get a Minkowski vacuum

but there are U(1)'s that only couple to the complex structure (stabilized) moduli,

So there is hope



Breaking the vacuum:

Since the BH drags the moduli out of their minimum, could it drag them so much that we'd tunnel into another vacuum? [D. Green, E. Silverstein, D. Starr '06]

Breaking the vacuum:

Since the BH drags the moduli out of their minimum, could it drag them so much that we'd tunnel into another vacuum? [D. Green, E. Silverstein, D. Starr '06]

Implications for fuzzballs?

Fuzzballs are resolutions of black hole horizons and singularities: the inside of the BH becomes a very stringy (and yet smooth) geometry.

In other words, they usually need UV physics to happen near the horizon (infinite distance in moduli space)

→ Fuzzballs **may** be in tension with scale separated vacua [Y. Li, '21]

Breaking the vacuum:

Since the BH drags the moduli out of their minimum, could it drag them so much that we'd tunnel into another vacuum? [D. Green, E. Silverstein, D. Starr '06]

Implications for fuzzballs?

Fuzzballs are resolutions of black hole horizons and singularities: the inside of the BH becomes a very stringy (and yet smooth) geometry.

In other words, they usually need UV physics to happen near the horizon (infinite distance in moduli space)

→ Fuzzballs **may** be in tension with scale separated vacua [Y. Li, '21]

WORK IN PROGRESS STAY TUNED!

